

Polarization of H₂O Masers in the presence of Velocity and Magnetic Field Gradients

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Abstract. Through polarization observations H₂O masers are excellent probes of magnetic fields in the maser region. Magnetic field strengths, such as those in the H₂O masers regions of the envelopes of late-type stars and star-forming regions, are typically determined using a direct relation between the field strength and the observed circular polarization. Here it is shown that velocity and magnetic field gradients along the maser have a significant effect on the field strengths obtained from circular polarization observations. Due to velocity gradients the actual magnetic field strength could be up to 100% higher than the field strength derived from the observations. Additionally, when a magnetic field gradient is present, the resulting circular polarization derived is caused predominantly by the average magnetic field in the unsaturated maser core. Measurements of the fractional linear polarization are not affected by velocity or magnetic field strength gradients, though changes in the magnetic field angle along the maser do quench the linear polarization intensity when the maser saturates.

Key words. masers – radiative transfer – polarization – magnetic fields

1. Introduction

Astrophysical H₂O masers, like other maser species such as OH and SiO, are excellent probes of magnetic fields in a variety of interesting regions. The magnetic field strength is derived from observations of maser polarization caused by Zeeman splitting. Polarization observations of the 22 GHz H₂O masers in the circumstellar envelopes (CSEs) around evolved stars have made important contributions to the understanding of the magnetic fields of evolved stars (Vlemmings, Diamond, & van Langevelde 2002; Vlemmings et al. 2005, hereafter V02 and V05a). H₂O maser observations have also provided information on the strength and structure of the magnetic field in star-forming regions (Fiebig & Guesten 1989; Sarma et al. 2001, 2002) and have provided upper limits for the field in the megamaser galaxy NGC 4258 (Modjaz et al. 2005).

However, the exact relationship between the observed maser polarization and the magnetic field strength is not straightforward, leading to uncertainties in the determined field strength. It was shown in Nedoluha & Watson (1992, hereafter NW92) that the influence of the different hyperfine components of the H₂O (6₁₆ – 5₂₃) rotational transition as well as of the degree of maser saturation is significant. The effect of maser saturation on the magnetic field strength determination can be as large as a factor of 2 (NW92). Additionally, velocity and magnetic field gradients might play an important role. While typical H₂O maser spot sizes are between $\sim 10^{12}$ and 10^{13} cm

(Reid & Moran 1981; Imai et al. 1997), the actual size of the maser region will be several times larger due to beaming (e.g. Goldreich & Keeley 1972). As a result the maser path length could be as long as $10^{14} - 10^{15}$ cm. Across such a path length, velocity and magnetic field gradients likely influence the observed polarization. In Vlemmings & van Langevelde (2005, hereafter V05b) it was shown that velocity gradients alter the shape of the total intensity line profile. From this analysis and from earlier observations (e.g. Richards et al. 1999) it is apparent that in CSEs velocity shifts of ~ 1 km s⁻¹ are common. In this paper the effects of velocity and magnetic field gradients along the maser path on the magnetic field strength determined from circular polarization measurements and on the observed fractional linear polarization are examined.

2. Background

Similar to the analysis in V02 and V05b, the equations of state for the populations of the upper (6₁₆) and lower (5₂₃) rotational levels of the three strongest hyperfine components ($F - F' = 7-6, 6-5$ and $5-4$) of the 22 GHz H₂O maser have been solved using the method described in NW92. The equation of state for the number density $n(F, a, \nu)$ of the upper energy levels (F) is

$$0 = \lambda_F(\nu_s) - (\Gamma + \Gamma_\nu)n(F, a, \nu) + R(F, F', a, b, \nu)(n(F', b, \nu) - n(F, a, \nu)) + \phi(\nu_s)\left(\frac{\Gamma_\nu}{\sum g_F}\right) \int d\nu \sum g_F n(F, a, \nu). \quad (1)$$

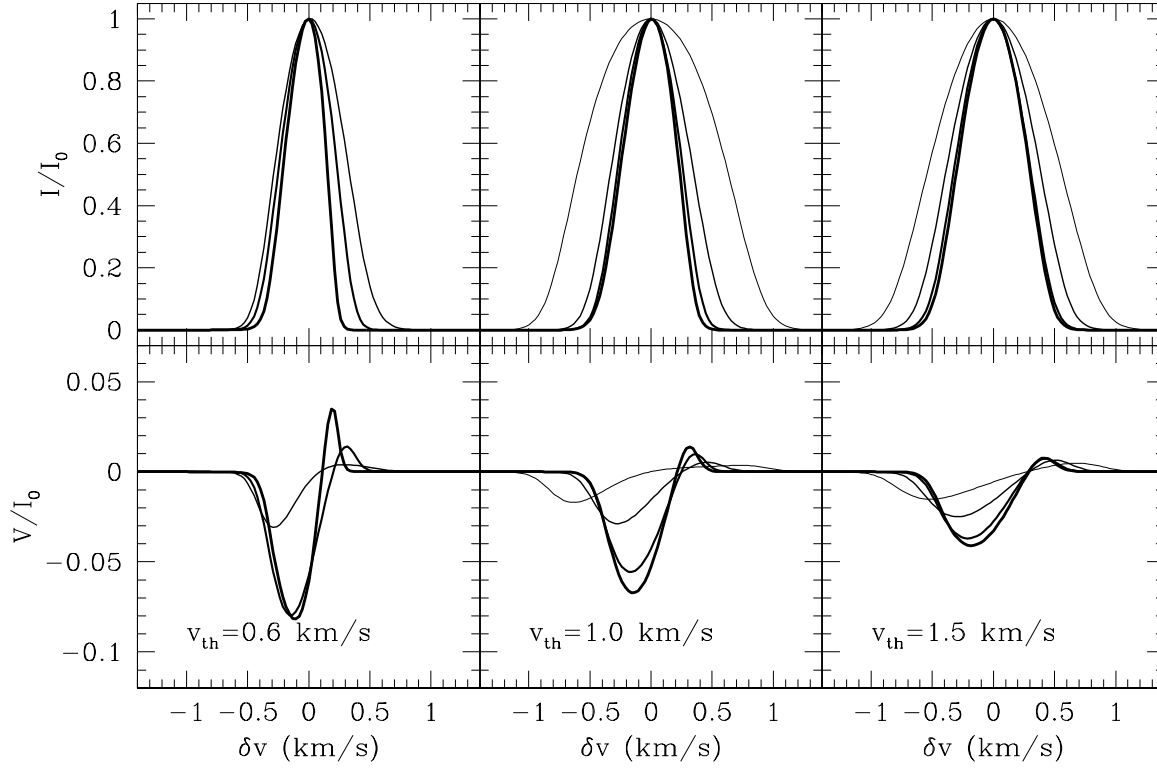


Fig. 1. Normalized total intensity (I) and circular polarization (V) spectra for $v_{\text{th}} = 0.6, 1.0$ and 1.5 km s^{-1} in the presence of a velocity gradient. From thick to thin the solid lines denote velocity shifts of $\Delta V_m = 0, 0.5, 1.0$ and 1.5 km s^{-1} respectively except for $v_{\text{th}} = 0.6 \text{ km s}^{-1}$ where $\Delta V_m = 1.5 \text{ km s}^{-1}$ is omitted.

The equation for the number density $n(F', b, v)$ of the lower levels (F') is similar but reversed. Here a and b denote the different magnetic sub-states of each hyperfine component. The population levels are solved as a function of molecular velocity v and at different positions along the maser propagation direction. The statistical weights are designated by g_F and g'_F , and are from Kukulich (1969). The pump rate $\lambda_F(v_s)$ is assumed to be the same for the different hyperfine components and has a Maxwellian distribution. The results of the calculations depend only on the ratio $(\lambda_F - \lambda'_F)/\lambda_F$, which is of the order of a few percent (Anderson & Watson 1993). Generally $v_s = v$, but when velocity gradients are introduced later, the molecular velocity of the maser pump v_s changes along the maser path. The rate for stimulated emission $R(F, F', a, b, v)$ (hereafter R) is calculated using the local maser intensity and the hyperfine interaction coefficients as described in NW92. Γ is the decay rate for the molecular excitations. The cross-relaxation rate Γ_v describes the reabsorption of previously emitted infrared pump photons trapped in optically thick transitions of the maser system, which generate a newly excited molecule at random, within a Maxwellian velocity distribution ($\phi(v_s)$). Elastic collisions between the maser molecules and intermixed H₂ molecules have not been included. The effects of these are described in detail in Elitzur (1990) and are unimportant for the H₂O masers discussed here.

The maser intensity, which is solved using the radiative transfer equations from NW92, influences the level populations through R . However, as it is itself dependent on the level populations determined with Eq. 1, the maser total intensity (I), linear polarization (Q) and circular polarization (V) are solved iteratively along the maser path. Assuming nearly one-dimensional maser propagation the beaming of the maser radiation is represented by the solid angle $\Delta\Omega$. The emerging maser fluxes will thus be represented in $T_b\Delta\Omega$, with T_b the brightness temperature. All calculations were performed with a brightness temperature for the radiation incident on the maser region of $T_b\Delta\Omega = 0.1 \text{ K sr}$. It was verified in NW92 and V02 that the results are insensitive to the chosen initial value. The calculations are performed for different thermal line widths (v_{th}) of the Maxwellian particle velocity distribution, where $v_{\text{th}} \approx 0.5(T/100)^{1/2} \text{ km s}^{-1}$ with T the temperature of the masing molecules in K. Nedoluha & Watson (1991) and NW92 have shown that when Eq. 1 is solved without including the contribution of the cross-relaxation rate Γ_v , the results for the line profiles, fractional polarizations and line width in the case of non-negligible Γ_v can be obtained by scaling the results with $[\Gamma + \Gamma_v]$. Following this result, the calculations in this paper are performed for $\Gamma = 1 \text{ s}^{-1}$ and $\Gamma_v = 0 \text{ s}^{-1}$. Thus, in all the results of this paper, the emerging brightness temperature $T_b\Delta\Omega$ can be replaced with $T_b\Delta\Omega \times [\Gamma + \Gamma_v]$ to obtain the results for different values of the decay and cross-relaxation rates.

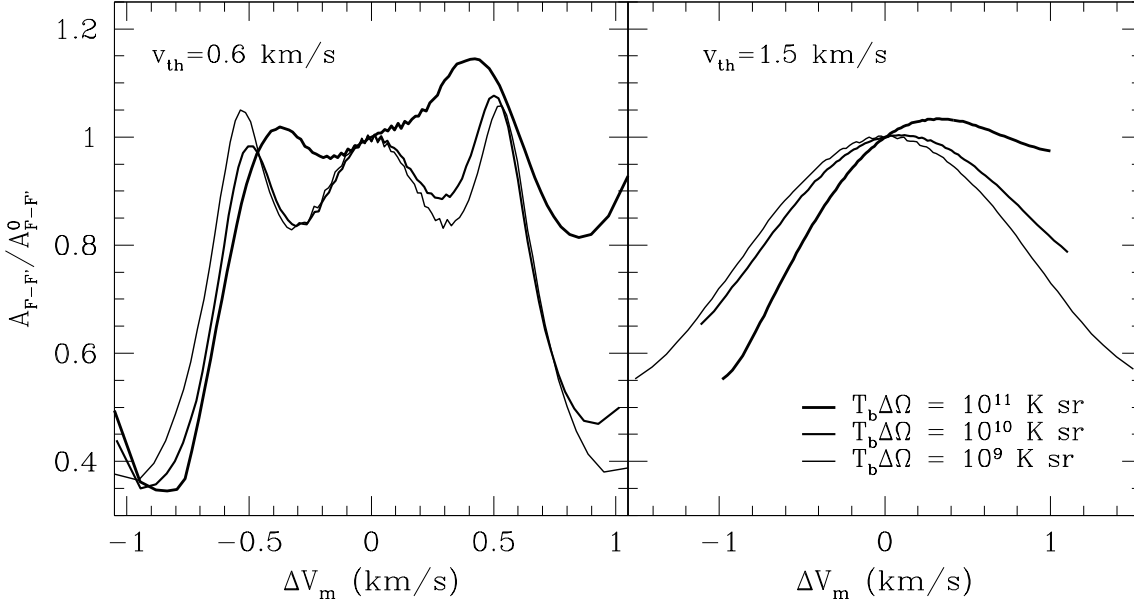


Fig. 2. The dependence of the magnetic field strength determined from circular polarization observation, represented by the $A_{F-F'}$ coefficient from Eq. 2, on the magnitude of the velocity gradient ΔV_m for two values of intrinsic thermal line width v_{th} . $A_{F-F'}$ is normalized by $A_{F-F'}^0$, which is $A_{F-F'}$ for $\Delta V_m = 0 \text{ km s}^{-1}$. The solid lines represent masers with different emerging brightness temperatures. Note that the noise ($< 1\%$) on the curves is due to the velocity resolution of our model spectra. The results are independent of the magnetic field angle θ .

To solve the equations of state it is assumed that the Zeeman frequency shift $g\Omega$ is much greater than R , Γ and Γ_v . Then the off-diagonal elements of the density matrix describing the molecular states are negligible, greatly simplifying the calculations. For a magnetic field strength of $\sim 1.0 \text{ G}$, $g\Omega$ of the dominant hyperfine transitions is $\sim 10^4 \text{ s}^{-1}$ while for $T_b \Delta\Omega < 10^{12} \text{ K sr}$ the rate for stimulated emission $R \leq 100 \text{ s}^{-1}$ (V02). Most of the calculations in this paper are performed for a magnetic field strength of $B = 1 \text{ G}$. For the highest $T_b \Delta\Omega$, the results will be valid for fields down to $\sim 10 \text{ mG}$, when $g\Omega$ becomes comparable to R . However, the low observed linear polarization of the 22 GHz H₂O masers (e.g. V02, V05a) indicates that $T_b \Delta\Omega$ is at most 10^{11} K sr and typically even lower. Thus, the results presented in this paper are valid for B well below 1 mG .

22 GHz H₂O masers begin to saturate when $R/\Gamma \gtrsim 1 \text{ s}^{-1}$ and approach full saturation when $R/\Gamma \gtrsim 100 \text{ s}^{-1}$. Thus, for our models with $\Gamma = 1 \text{ s}^{-1}$, saturation starts at $T_b \Delta\Omega \gtrsim 10^{10} \text{ K sr}$ and full saturation occurs for $T_b \Delta\Omega \gtrsim 10^{12} \text{ K sr}$. The maser spectral line profile (through rebroadening) and the polarization properties of the maser changes at $R = (\Gamma + \Gamma_v)$ (e.g. NW92). While typically for the 22 GHz H₂O masers this is close to when saturation occurs, a large cross-relaxation rate means the changes will only occur when the maser has already become partly saturated. As Γ_v for the astrophysical H₂O masers is at most $\approx 5 \text{ s}^{-1}$ for H₂O masers at $T \sim 1000 \text{ K}$ and is less than that lower temperatures (Anderson & Watson 1993), $\Gamma_v \approx \Gamma$ is assumed in this paper when discussing unsaturated and saturated masers. It is worth noting that H₂O maser

observations indicate that in actuality Γ is between one and two orders of magnitude less than Γ_v (Anderson & Watson 1993). However, as our results scale linearly with $(\Gamma + \Gamma_v)$ this does not affect the conclusions in this paper.

3. Radiative Transfer Modeling Results

3.1. Circular Polarization

As shown in NW92 and V02, the V-spectra produced from the radiative transfer equations are not anti-symmetric. They are also not always directly proportional to the derivative of the I-spectrum due to hyperfine interaction as is often assumed. However, as was discussed in V02, it is impossible to directly observe the intrinsic shape of the V-spectrum due to instrumental effects and necessary data calibration steps. Still, the percentage of circular polarization P_V can be related to the magnetic field strength using

$$P_V = (V_{\max} - V_{\min})/I_{\max} = 2 \cdot A_{F-F'} \cdot B_{[\text{Gauss}]} \cos\theta / \Delta v_L [\text{km s}^{-1}]. \quad (2)$$

Here V_{\max} and V_{\min} are the minimum and maximum of the V-spectrum. I_{\max} and Δv_L are the peak flux and the full width half-maximum (FWHM) of the I-spectrum respectively. B is the magnetic field strength at an angle θ to the maser propagation direction. While in most analyses of maser circular polarization, the coefficient $A_{F-F'}$ is taken to be a fixed value (e.g. Fiebig & Guesten 1989), it was shown in NW92 to depend on the level of maser saturation and the intrinsic thermal velocity v_{th} of the maser. Figure 6 of V02 shows $A_{F-F'}$

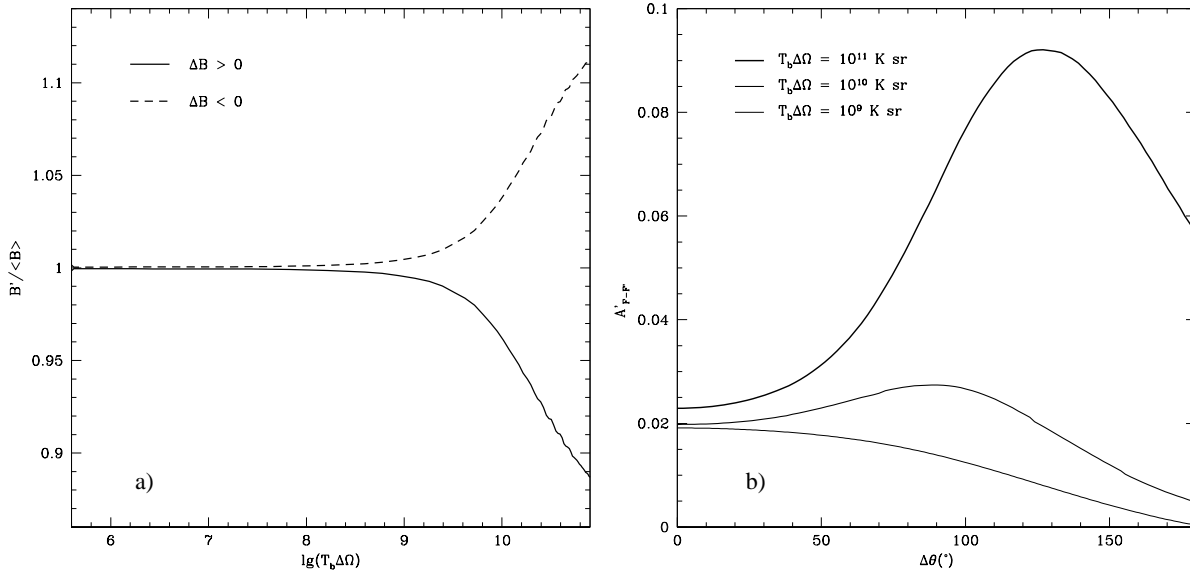


Fig. 3. a) Derived magnetic field strength B' over the average field strength along the maser $\langle B \rangle$ as a function of $T_b \Delta\Omega$. The solid and dashed lines correspond to positive and negative magnetic field gradients respectively. b) $A'_{F-F'}$ as a function of $\Delta\theta$ for different values of emerging brightness temperature $T_b \Delta\Omega$. Calculations were performed for $v_{th} = 0.8 \text{ km s}^{-1}$.

for different values of v_{th} as a function of emerging maser brightness temperature $T_b \Delta\Omega$. For maser brightness temperatures $T_b \Delta\Omega > 10^9 \text{ K sr}$ it was shown in NW92 that the $\cos\theta$ dependence of Eq. 2 breaks down effectively also introducing a dependence on θ to $A_{F-F'}$. This was later shown in more detail in Watson & Wyld (2001) for masing involving angular momentum $J = 1-0$ and $J = 2-1$ transitions. In V02, figure 7 shows the derived magnetic field strength dependence on θ to $A_{F-F'}$ for the 22 GHz $J = 6-5$ transition.

3.1.1. velocity gradients

First, a velocity gradient is introduced by uniformly shifting, in Eq. 1, the velocity v_s of the pump rate λ_F for each integration step along the maser path. By varying the amount of shift, a velocity difference ΔV_m is created between the start and end of the maser amplification path of up to $\sim 1.5 \text{ km s}^{-1}$. This corresponds to a velocity gradient of $\Delta V_m / S \text{ km s}^{-1} \text{ m}^{-1}$, where S is the length of the maser path. As S depends on the exact values of the pumping rate λ_F , while our results, expressed in emerging brightness temperature are only dependent on the ratio of pumping rates as addressed above, any mention of the velocity gradient will henceforth be referring to the value of ΔV_m . In Nedoluha & Watson (1988) it was shown that the maser splits into several distinctly separate narrow features when ΔV_m becomes a few times v_{th} , each of the features becoming an independent maser. The calculations in this paper are therefore limited to values of ΔV_m up to $\sim 1.5 \text{ km s}^{-1}$.

Fig. 1 shows the effect of a velocity gradient on the total intensity and circular polarization profiles for three different intrinsic thermal velocities v_{th} , produced for the same magnetic field strength (1 G). While for low ΔV_m , the effect of the velocity gradient on the shape of the I-profile is small, the effect

on the V-spectrum can still clearly be seen. To examine how a velocity gradient influences the magnetic field determination from the circular polarization measurements, the normalized $A_{F-F'}$ coefficient from Eq. 2 is plotted in Fig. 2 as a function of velocity gradient for $v_{th} = 0.6$ and 1.5 km s^{-1} and emerging brightness temperatures $T_b \Delta\Omega = 10^9, 10^{10}$ and 10^{11} K sr . Because of the normalization, the results are independent of the magnetic field angle θ . For $v_{th} = 0.6 \text{ km s}^{-1}$ the effect of the hyperfine components is clear. For $T_b \Delta\Omega = 10^9 \text{ K sr}$ none of the hyperfine components are becoming saturated yet, thus $A_{F-F'}$ is mostly symmetric around $\Delta V_m = 0 \text{ km s}^{-1}$. The symmetry breaks down when the hyperfine lines are slowly starting to saturate for $T_b \Delta\Omega \gtrsim 10^{10} \text{ K sr}$, as maser growth is somewhat inhibited for velocity gradients in the direction of the weaker hyperfine components ($F - F' = 6 - 5$ and $5 - 4$), while a gradient in the opposite direction allows for stronger maser amplification. As noted above this effect can be postponed further into the saturated regime when Γ_v increased. It can be seen in the figure, that $A_{F-F'}$ typically decreases for increasing $|\Delta V_m|$, implying that in the presence of velocity gradients the magnetic field strength B' derived with Eq. 2 underestimates the true B . For small v_{th} the field strength can sometimes be underestimated by more than a factor of 2. However, for a narrow range in ΔV_m , $A_{F-F'}$ is actually enhanced. In those cases B could actually be overestimated by up to $\sim 20\%$. This effect is largest for the more saturated masers.

3.1.2. magnetic field gradients

There are different ways a magnetic field gradient along the maser amplification path can occur. The actual field strength B can change or the angle θ between the field and the maser propagation direction can vary. Similar to ΔV_m , differences $\Delta\theta$

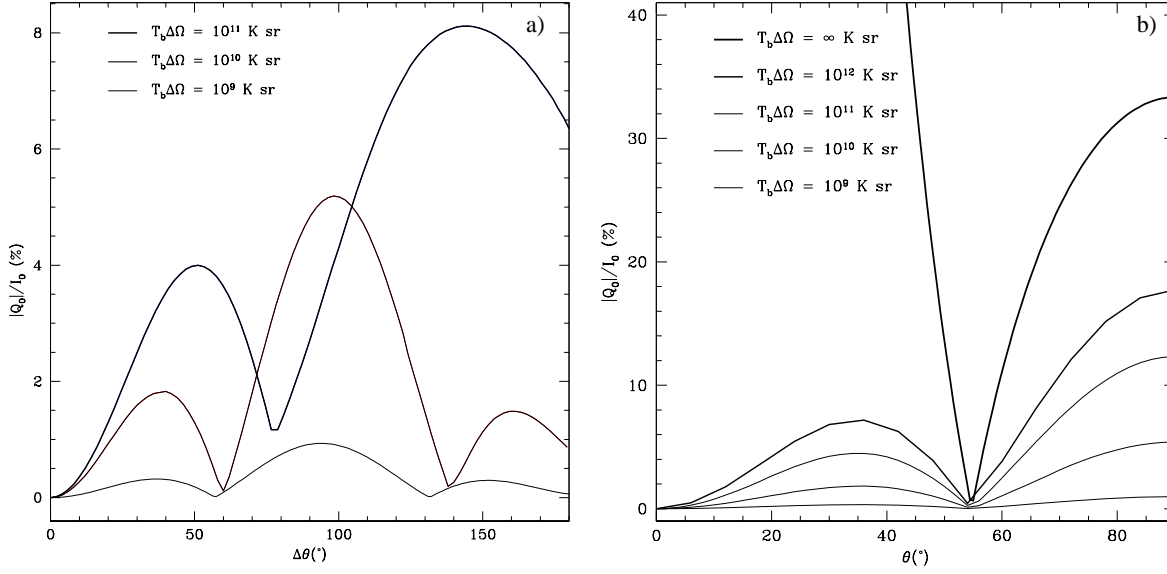


Fig. 4. a) Magnetic field angle gradient $\Delta\theta$ vs. the fractional linear polarization for three different stages of maser saturation and $\theta_0 = 0^\circ$. b) The angle θ between the maser propagation direction and the magnetic field vs. the fractional linear polarization for different values of emerging maser brightness. The thick solid line denotes the theoretical limit from Goldreich et al. (1973) for a completely saturated maser.

and ΔB are introduced between the start and end of the maser path and their influence on the maser polarization is examined. Fig. 3a shows the ratio between the magnetic field B' determined with Eq. 2 and the average magnetic field $\langle B \rangle$ along the maser as a function of $T_b\Delta\Omega$ when a gradient ΔB is included. The shape of this function does not depend on the actual value of ΔB or on the intrinsic thermal line width v_{th} . While the maser is unsaturated, or more precisely while $R < \Gamma_\nu$, B' is equal to $\langle B \rangle$. However, when the maser saturates, B' will be dominated by the average magnetic field in the unsaturated maser region.

Fig. 3b shows the change in $A'_{F-F'}$ versus $\Delta\theta$ for different $T_b\Delta\Omega$. The angle θ_0 , the value of θ at the start of maser amplification is taken to be 0° . Here $A'_{F-F'} = A_{F-F'} \cos\theta$, as the actual angle θ cannot be specified due to the inclusion of the gradient $\Delta\theta$. Comparing the unsaturated maser ($T_b\Delta\Omega = 10^9$ K sr) with the partly saturated maser ($T_b\Delta\Omega = 10^{11}$ K sr) it is again apparent that, as seen in Fig. 3a, the unsaturated maser regime provides the dominant contribution to the magnetic field strength determined with Eq. 2. Comparing Fig. 3b with figure 7 from V02 indicates that the angle average $\langle\theta\rangle$ in the unsaturated maser region determines the observed circular polarization.

3.2. Linear Polarization

The linear polarization in the presence of a magnetic field was also discussed in NW92. The fractional linear polarization is not found to be affected when including a velocity gradient ΔV_m or magnetic field gradient ΔB in the calculations. Only when including a gradient $\Delta\theta$ does the linear polarization change. In Fig. 4a the fractional linear polarization $|Q_0|/I_0$ is shown as a function of $\Delta\theta$. This can be compared with $|Q_0|/I_0$ as a function of the angle between the magnetic field and the

maser propagation direction θ shown in Fig. 4b. This relationship, in the limiting case of a completely saturated maser, was solved in Goldreich et al. (1973). In the case of an unsaturated maser ($T_b\Delta\Omega \lesssim 10^{10}$ K sr), the relationship between $|Q_0|/I_0$ and the magnetic field angle gradient $\Delta\theta$ is very similar to the relationship between $|Q_0|/I_0$ and θ , while when the saturation level increases the linear polarization is quenched and shifted. This indicates that the observed fractional linear polarization is determined by θ at the end of the unsaturated maser regime, where the largest amplification occurs, and not by the average $\langle\theta\rangle$ over the unsaturated maser core.

4. Conclusions

Using a maser radiative transfer code which includes the three strongest hyperfine components of the 22 GHz H₂O $6_{16} - 5_{23}$ rotational transition and their magnetic sub-states, the effects of velocity and magnetic field gradients on the determination of magnetic fields from maser polarization observations have been calculated. It was shown that, due to velocity gradients, the magnetic field strength determined on H₂O masers in CSEs (V02, V05a), where $v_{th} \approx 0.8-1.0$ km s⁻¹ and where velocity gradients of ~ 1.0 km s⁻¹ have been observed (V05b), can be underestimated by up to a factor of 2. However, for maser features with smaller velocity gradients there will be cases where B is actually overestimated by $\sim 10-20\%$. The magnetic fields determined on the H₂O masers in star-formation regions (Fiebig & Guesten 1989; Sarma et al. 2001, 2002), which have higher thermal line widths ($v_{th} \approx 1.5$ km s⁻¹) can be underestimated by $\sim 20-40\%$. These uncertainties can only be overcome by including a consistent model for the velocity gradients

throughout the H₂O maser source when deriving the magnetic field strengths derived from circular polarization observations.

When there is a magnetic field gradients along the maser amplification path, B' determined from polarization observations of an unsaturated maser (more formally when $R \lesssim \Gamma_\nu$), is typically equal to the average field along the maser. When the maser is saturated, the B' is dominated by the average magnetic field strength or the average magnetic field angle in the unsaturated maser core. The fractional linear polarization on the other hand, is not affected by velocity or magnetic field gradients. Only a gradient in the angle between the magnetic field and the maser propagation axis alters the linear polarization. In contrast to the circular polarization, the fractional linear polarization is mainly determined by the magnetic field angle in the part of the unsaturated maser where the largest amplification occurs and not by the angle averaged over the entire unsaturated maser region.

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